

# Emulators for Inverse Problems in Dense Matter Physics

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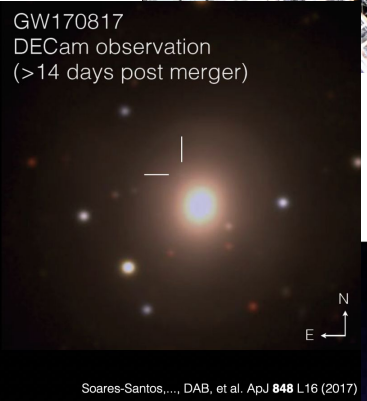
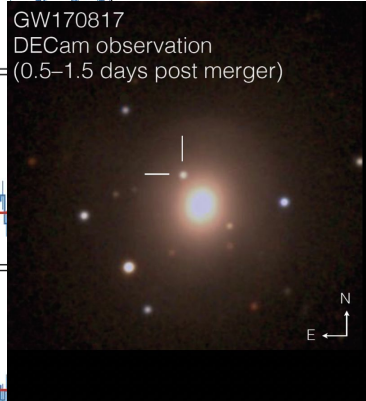
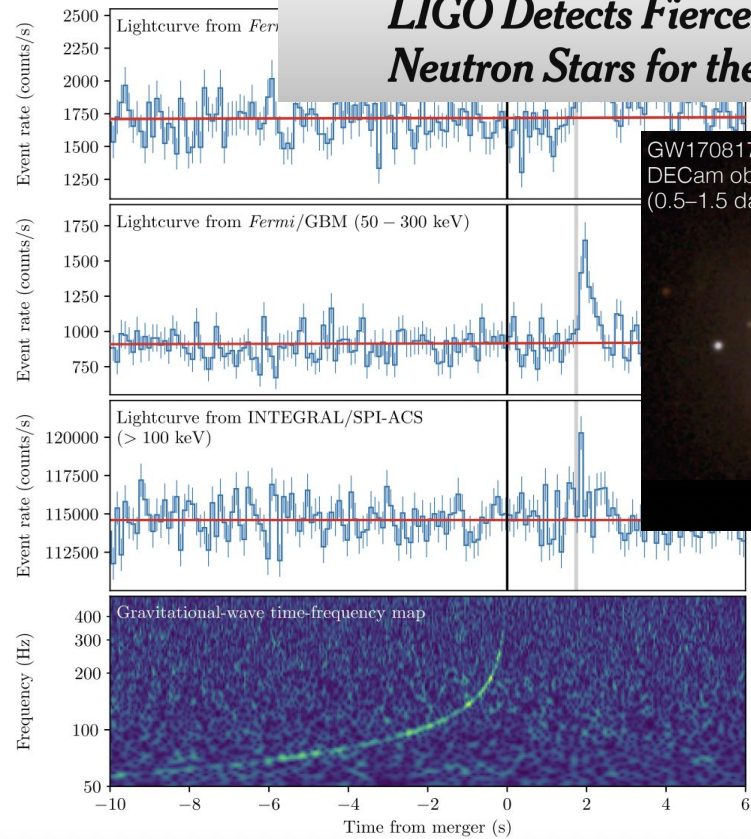
07/10/2024, INT Workshop: Inverse Problems and Uncertainty Quantification in Nuclear Physics



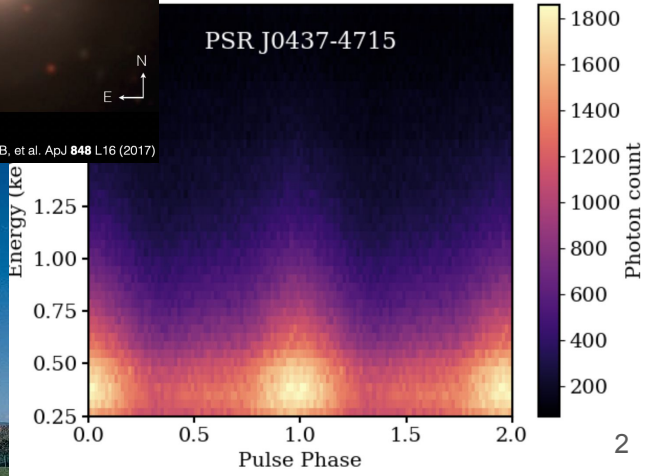
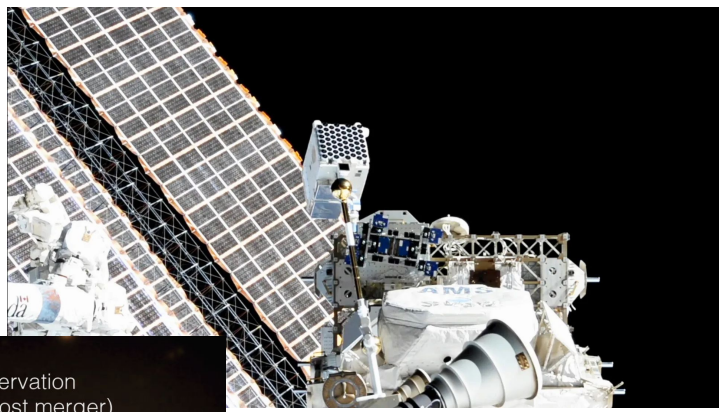
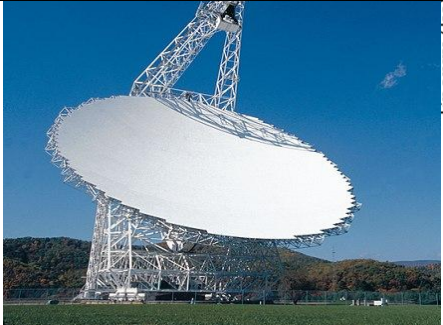
# An explosion of NS observations!

The New York Times

## LIGO Detects Fierce Collision of Neutron Stars for the First Time

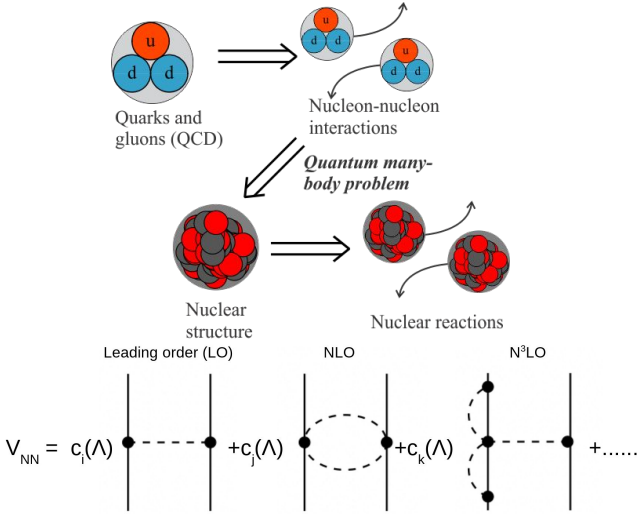


Soares-Santos, ..., DAB, et al. ApJ 848 L16 (2017)



# Dense matter physics in a nutshell

## Model for interaction between particles

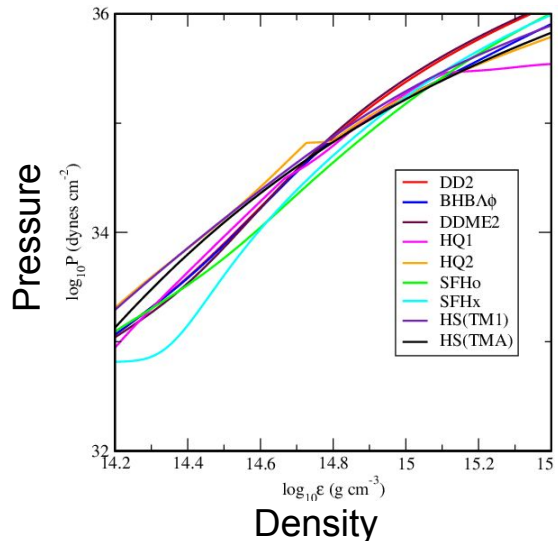
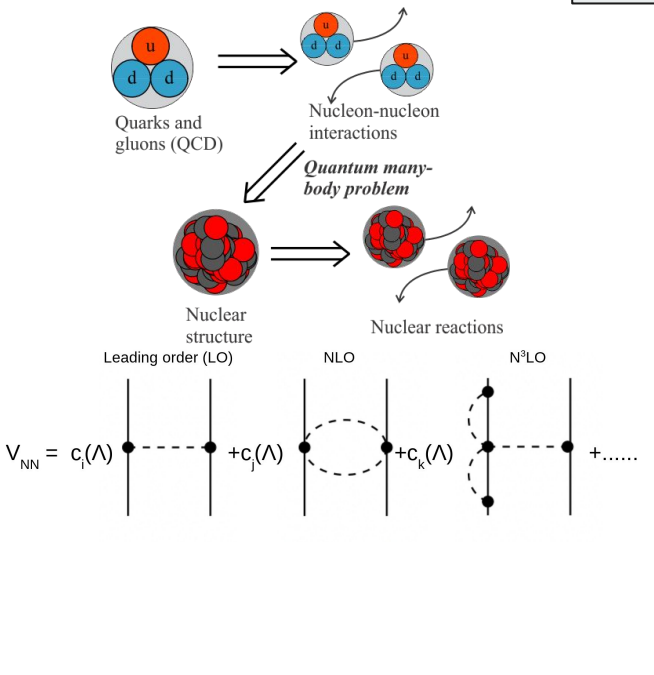


# Dense matter physics in a nutshell

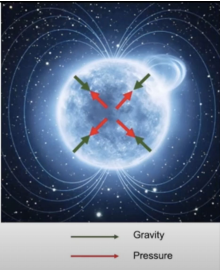
Model for interaction between particles

$$H|\psi\rangle = E|\psi\rangle$$

The Equation of State



# Dense matter physics in a nutshell



Tolman–Oppenheimer–Volkoff (TOV) equation

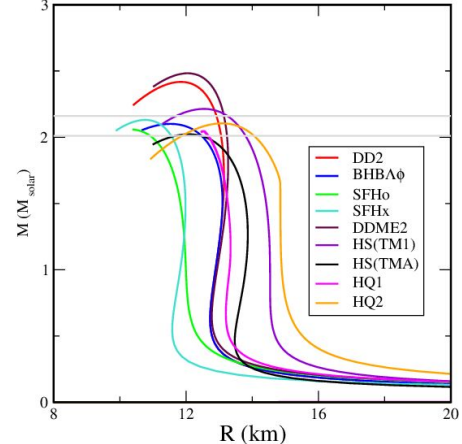
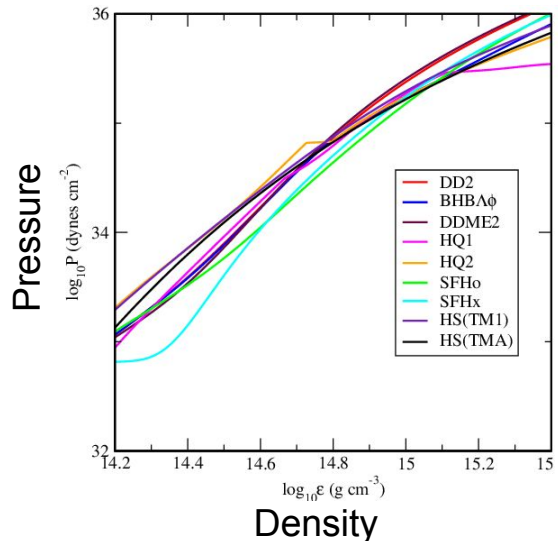
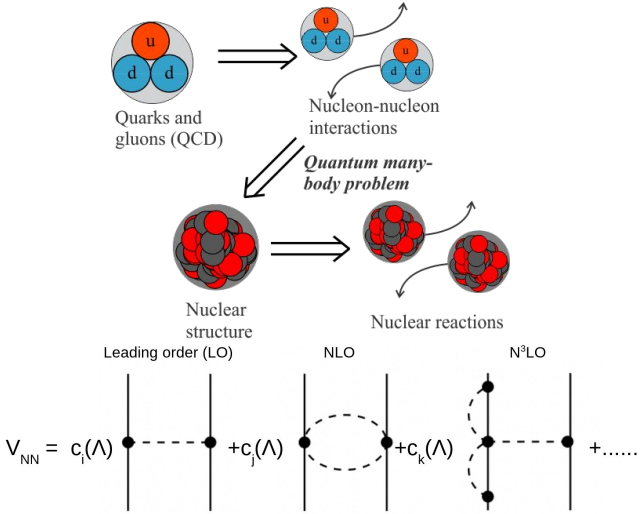
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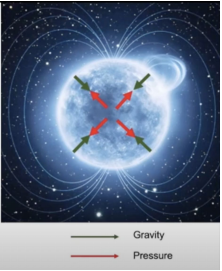
$$G^{\mu\nu} = \kappa T^{\mu\nu}$$

Neutron star observables



# Dense matter physics in a nutshell

Bayesian inference requires  $\sim 10^7$  model evaluations



Tolman–Oppenheimer–Volkoff (TOV) equation

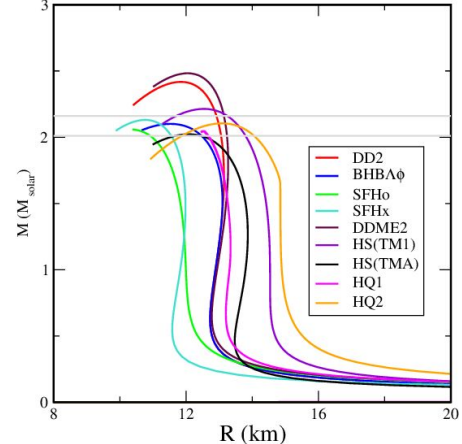
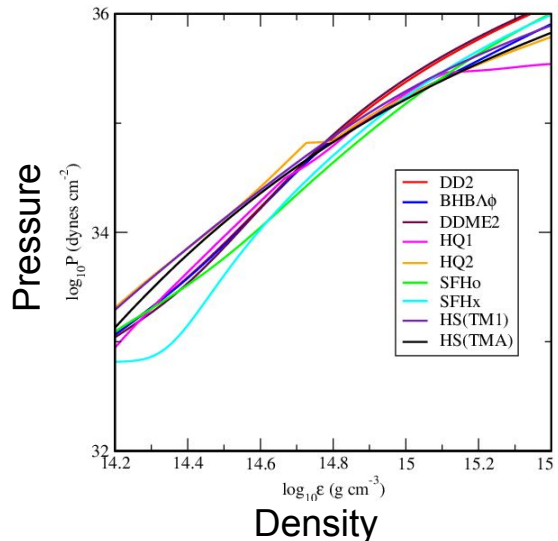
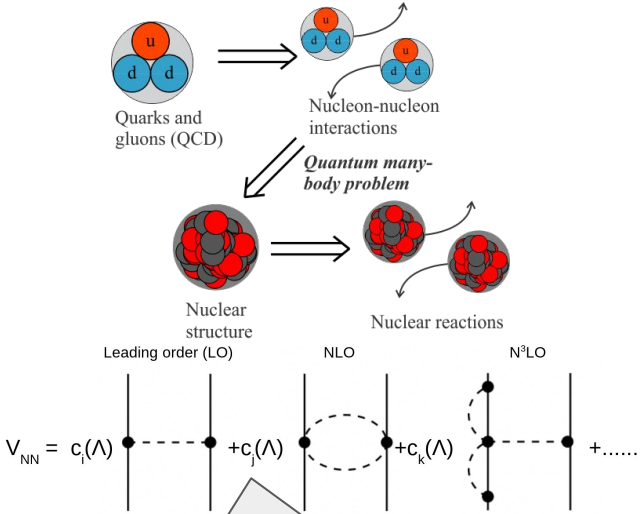
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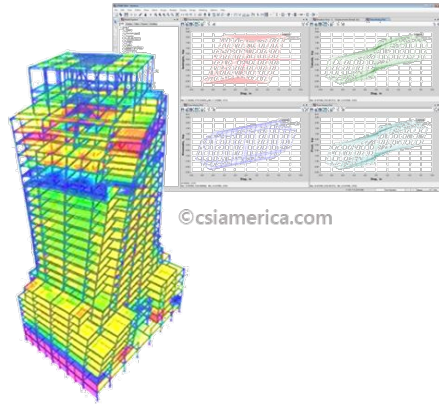
Bayesian Inference methods



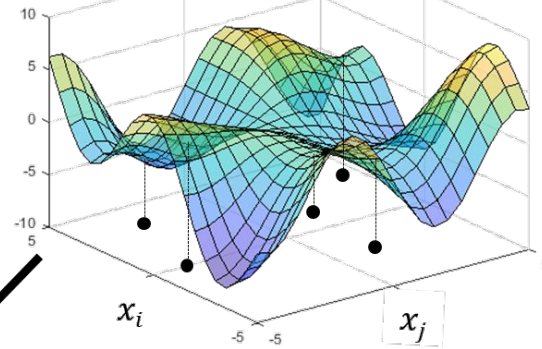
# The solution: Use emulators to accelerate calculations

Emulators mimic the behaviour of the full-scale model at a small fraction of its computational cost

## Simulation model

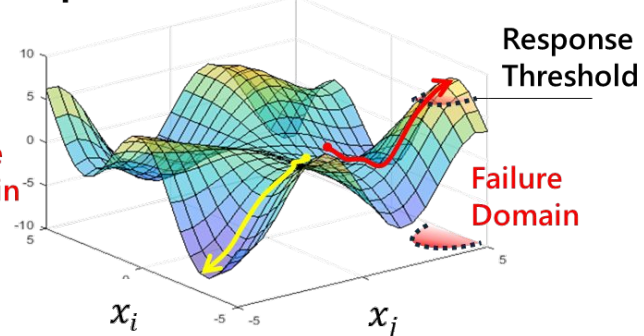
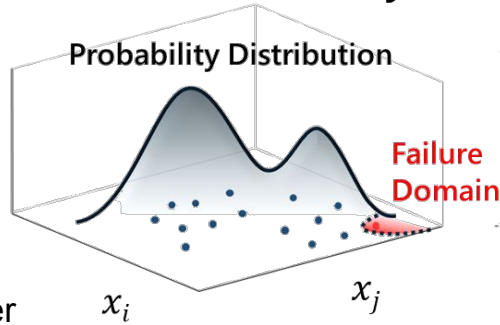


## Surrogate model



Training

## UQ Analysis / Optimization



# Overall strategy

Quantum many-body methods  
 ~1 million CPU hours  
 (scarce training data)

TOV equations  
 ~5 seconds

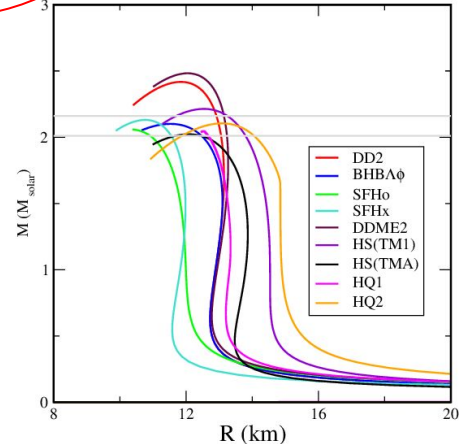
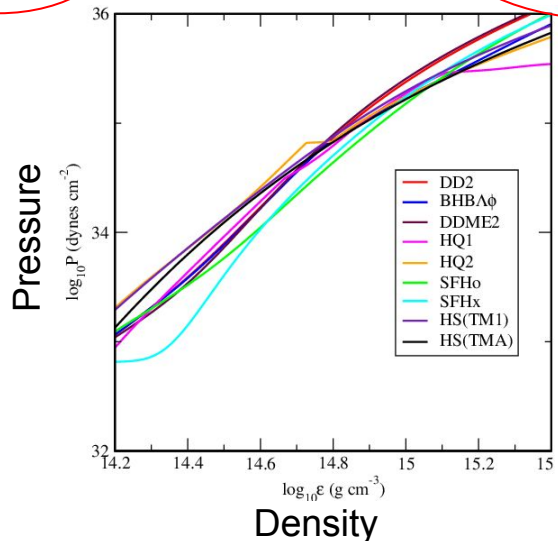
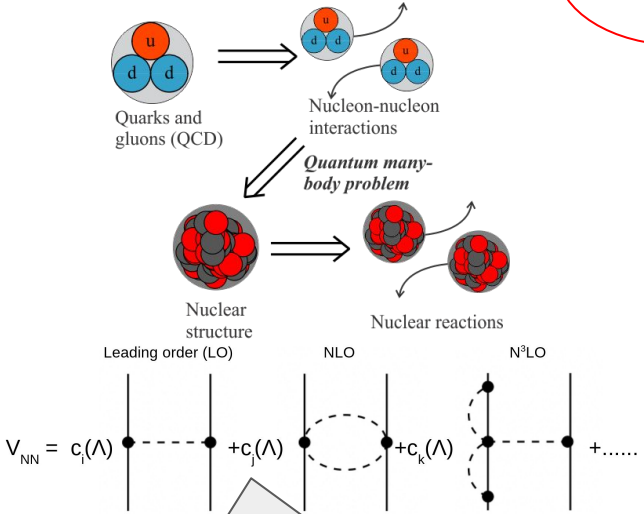
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Bayesian Inference methods



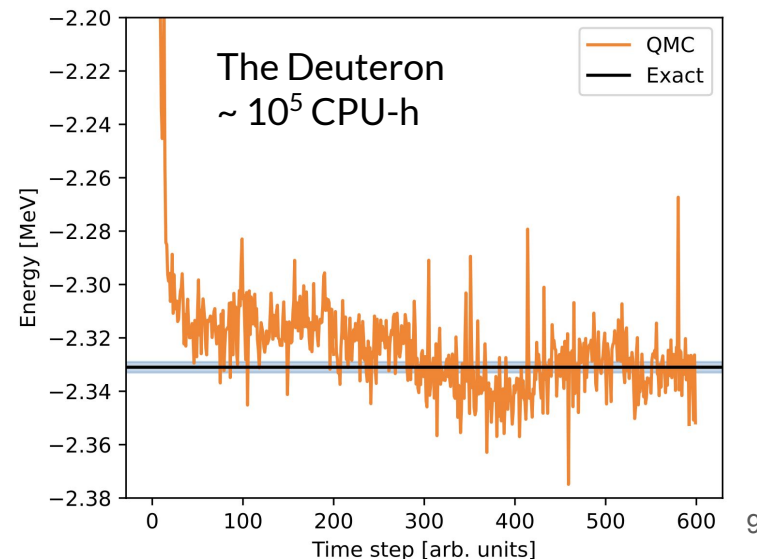
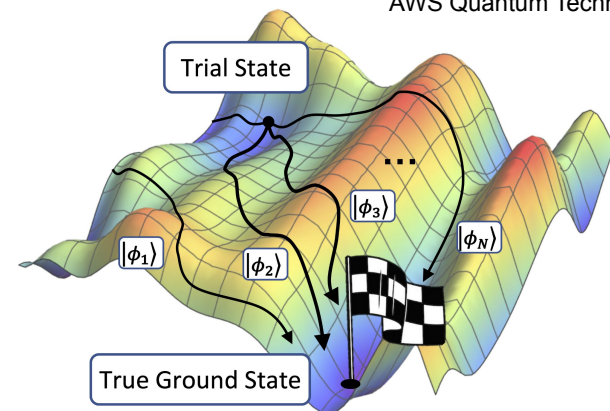
# Quantum Monte Carlo in a nutshell

$$\lim_{\tau \rightarrow \infty} e^{-H\tau} |\Psi_T\rangle \rightarrow |\Psi_0\rangle$$

Trial wavefunction
True ground state

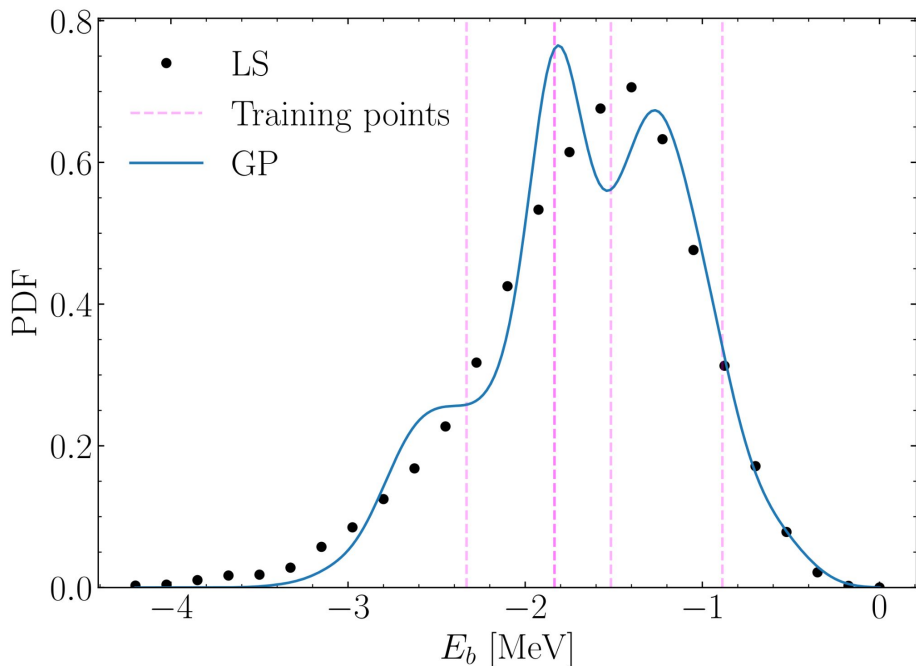
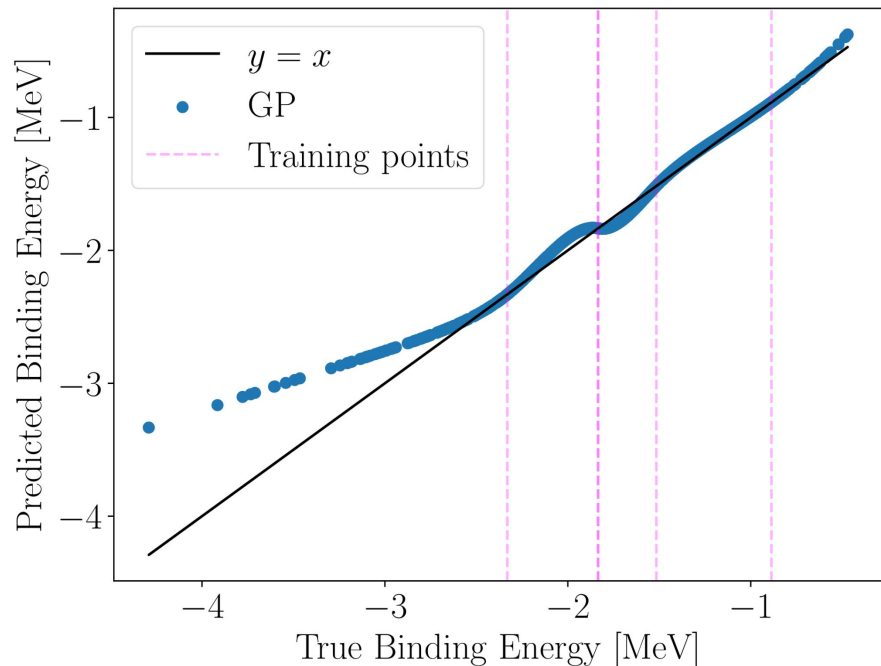
$H(\mathbf{a})$  is the Hamiltonian

- Virtually exact method for strongly interacting many-body systems
- First step is the preparation of a trial wavefunction, i.e. our best guess for the true ground state
- The trial state is evolved in imaginary time. This is mathematically equivalent to the diffusion problem
- At infinite imaginary time, the system ‘cools’ to its true ground state



# Emulators with scarce data: how about traditional ML?

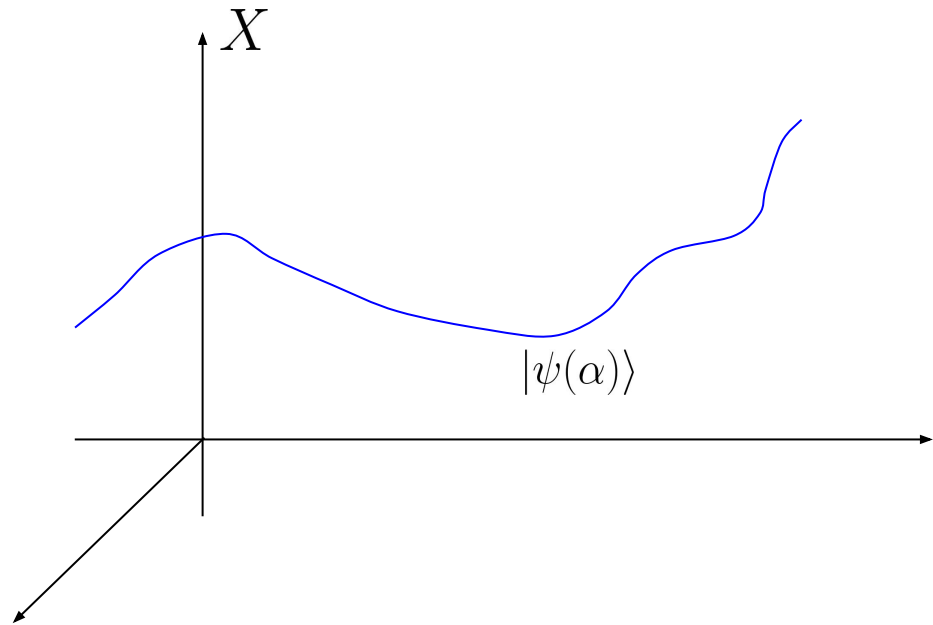
- Goal: Build accurate surrogate models for QMC with  $\sim 5 - 10$  training points
- The GP fails to accurately interpolate and extrapolate between training points



# Emulators with scarce data: Reduced basis methods

- Start with the full Schrödinger equation

$$H|\psi\rangle = E|\psi\rangle$$

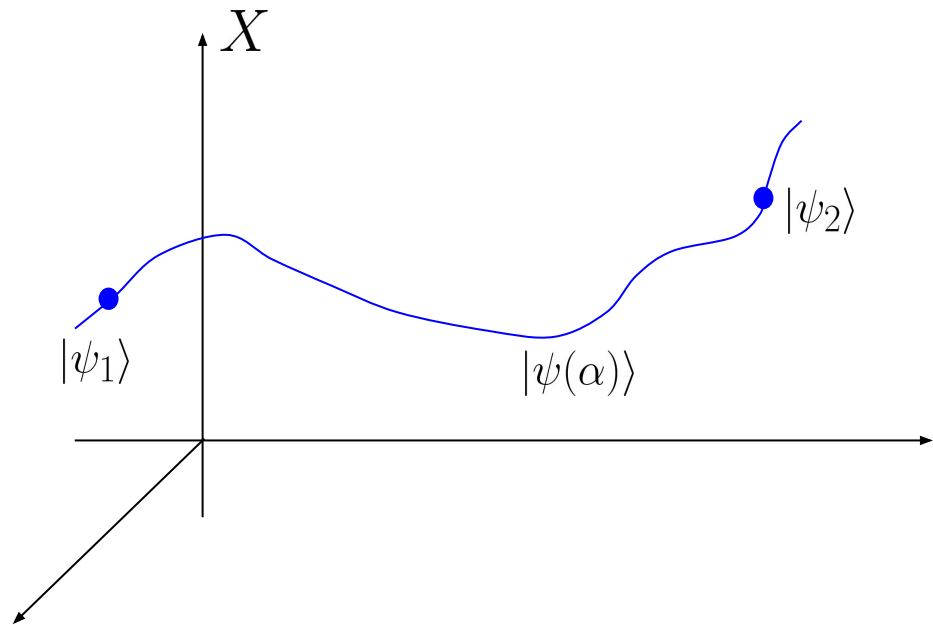


# Emulators with scarce data: Reduced basis methods

- Start with the full Schrödinger equation

$$H|\psi\rangle = E|\psi\rangle$$

- Compute N training functions or ‘snapshots’  $|\psi_j\rangle$



# Emulators with scarce data: Reduced basis methods

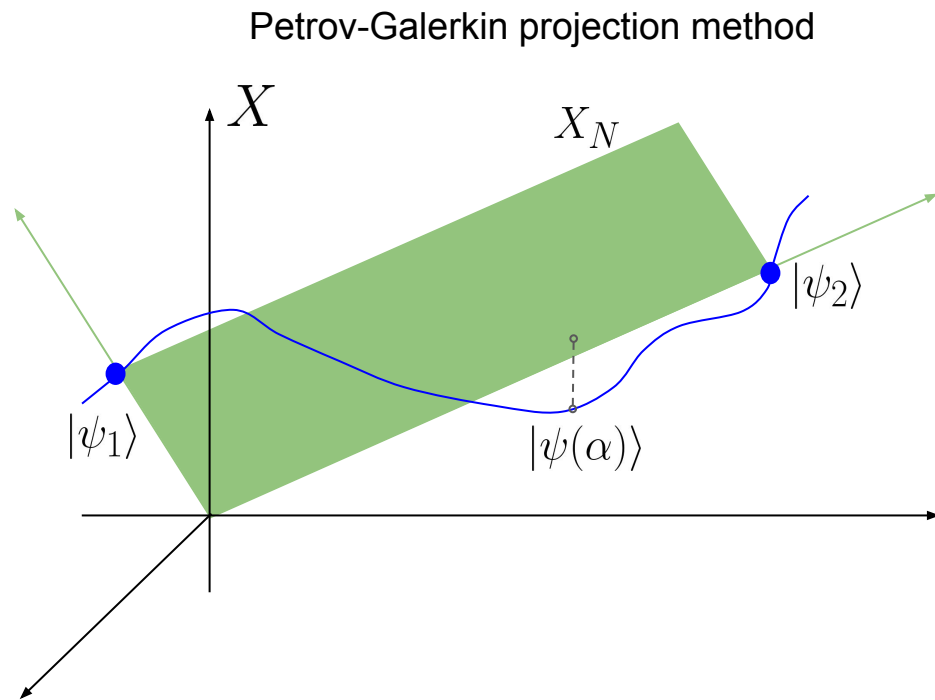
- Start with the full Schrödinger equation

$$H|\psi\rangle = E|\psi\rangle$$

- Compute  $N$  training functions or 'snapshots'  $|\psi_j\rangle$
- Project the Hamiltonian into the reduced space spanned by  $|\psi_j\rangle$

Mathematically, this corresponds to computing the matrix

$$M_{ij} \equiv \langle \psi_i | H | \psi_j \rangle$$



# Emulators with scarce data: Reduced basis methods

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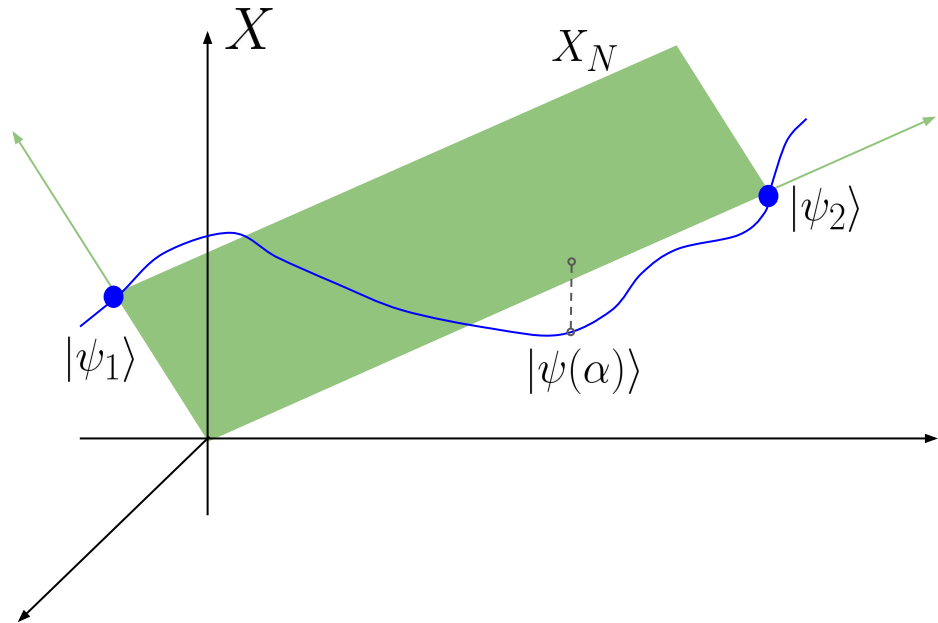
Mathematically, this corresponds to computing the matrix

$$M_{ij} \equiv \langle \psi_i | H | \psi_j \rangle$$

- For QMC,  $M_{ij}$  is dominated by stochastic noise and cannot be calculated. We therefore implemented a Petrov-Galerkin projection method for this problem:

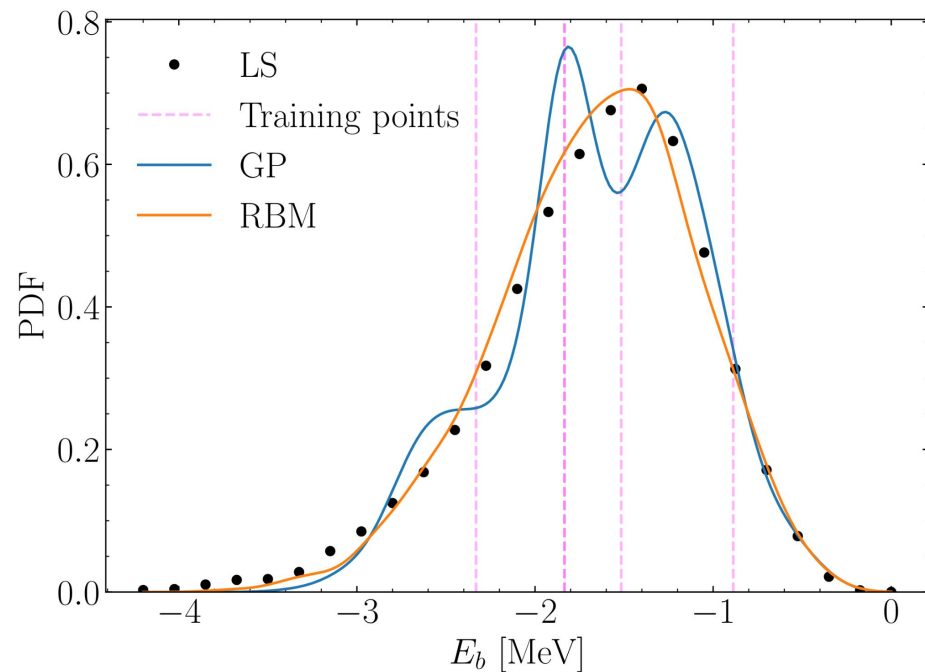
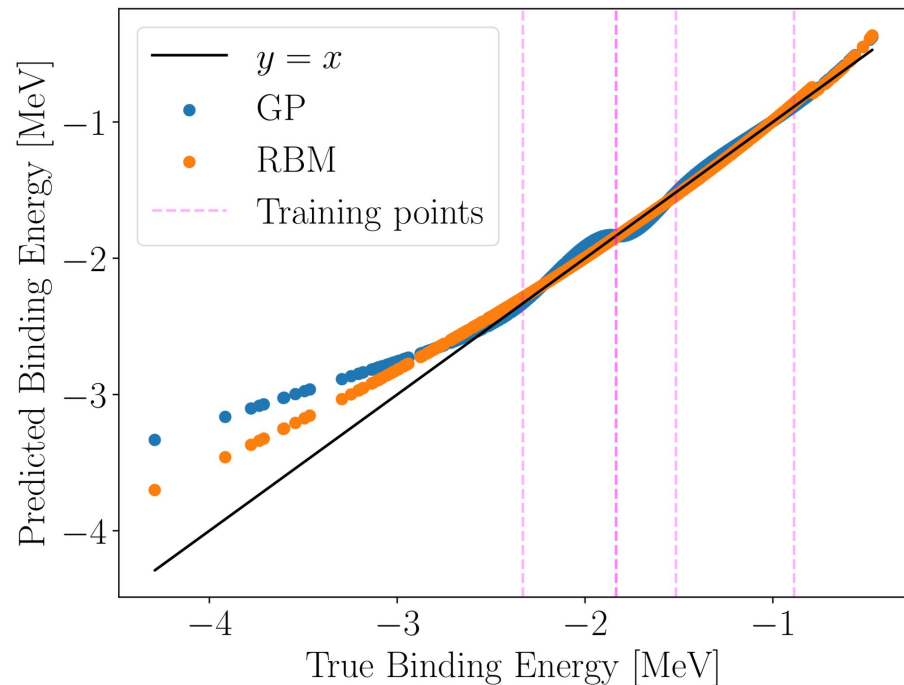
$$\tilde{M}_{ij} = \langle \psi_i^T | H | \psi_j \rangle$$

Petrov-Galerkin projection method



# Emulators with scarce data: Reduced basis methods

- The RBM outperforms the GP
- The RBM is capable of interpolating but fails to extrapolate away from training points



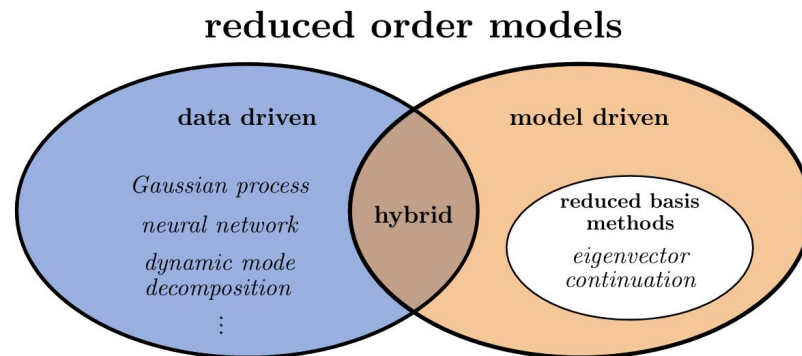


# Emulators with scarce data: Hybrid models

- Combine elements of RBMs with data-driven emulators
- We employ the recently proposed parametric matrix models

$$M(\vec{\alpha}) = M_0 + \sum_i \alpha_i M_i$$

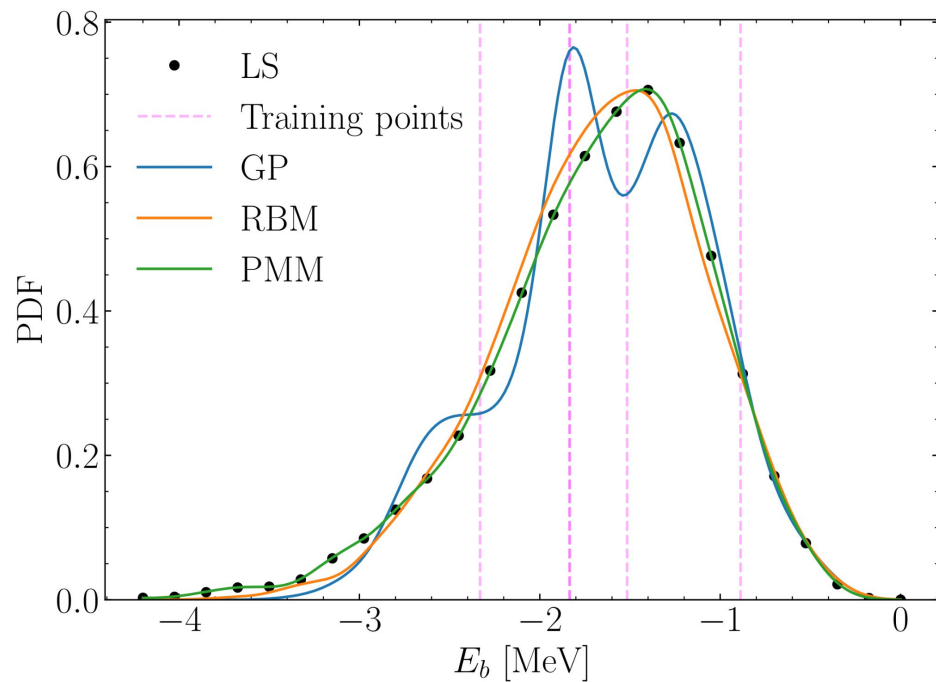
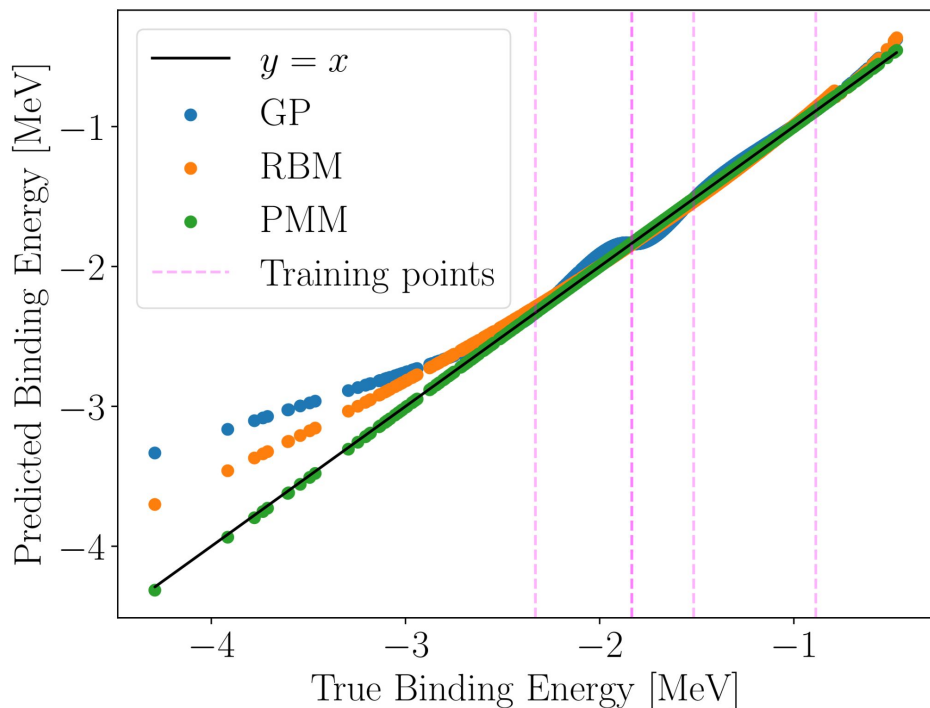
- The form of the reduced subspace matrix is inspired by RBMs. However, we do not directly compute the projections, i.e. we do not compute the subspace matrix elements
- Instead they are learned in some manner from the data



Duguet et al., arXiv:2310.19419

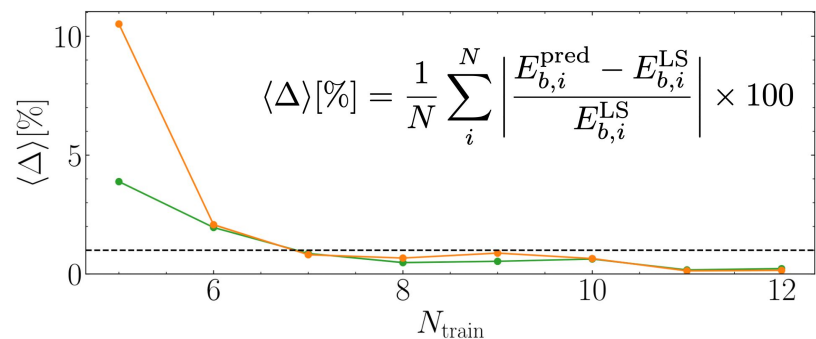
# Emulators with scarce data: Hybrid models

- The PMM outperforms both the GP and RBM
- It interpolates well but also gives excellent results for extrapolation!

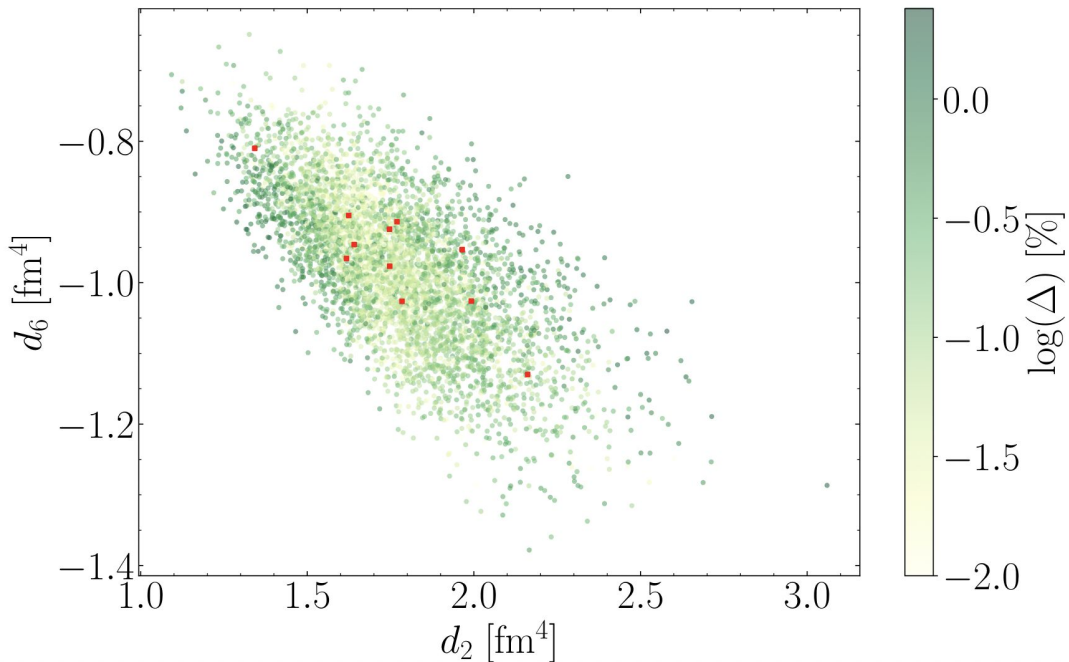


# Emulators with scarce data: Hybrid models

- The PMM outperforms both the GP and RBM.
- It interpolates well but also gives excellent results for extrapolation!



- We found that our PMM generalises very well to at least 4 - 5 dimensional parameter spaces.



# Overall strategy

Speed-up factor  $\sim 10^7$   
 Uncertainty  $\sim 0.1\%$

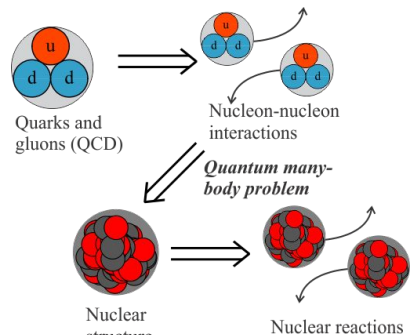
TOV equations  
 $\sim 5$  seconds



Model for interaction between particles  $\rightarrow$  The Equation of State  $\rightarrow$  Neutron star observables

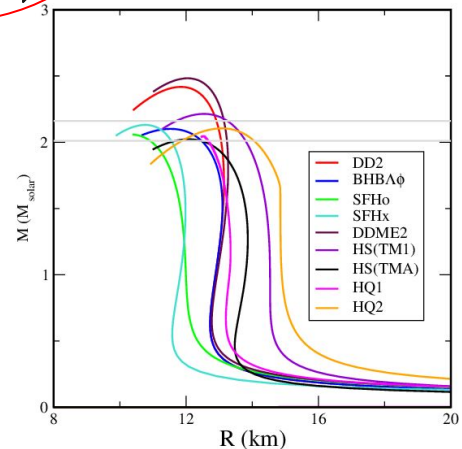
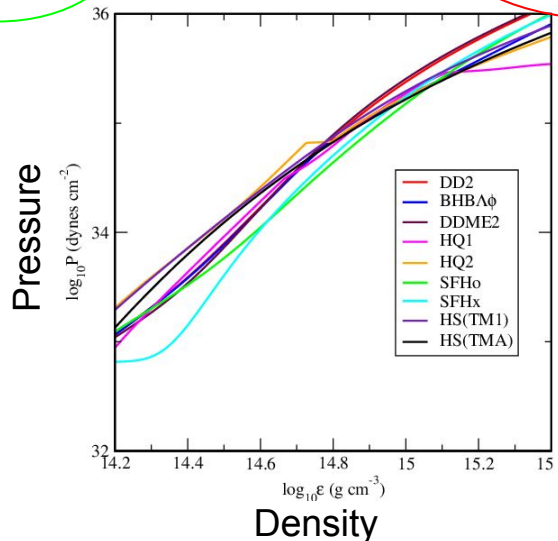
$$H|\psi\rangle = E|\psi\rangle$$

$$G^{\mu\nu} = \kappa T^{\mu\nu}$$



$$V_{NN} = c_1(\Lambda) + c_2(\Lambda) + c_3(\Lambda) + c_4(\Lambda) + \dots$$

Leading order (LO) NLO N<sup>3</sup>LO

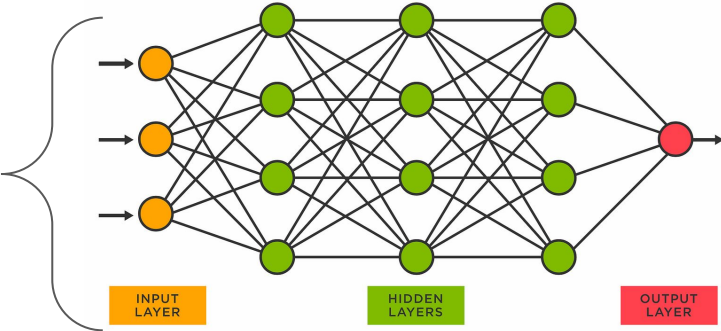


Bayesian Inference methods

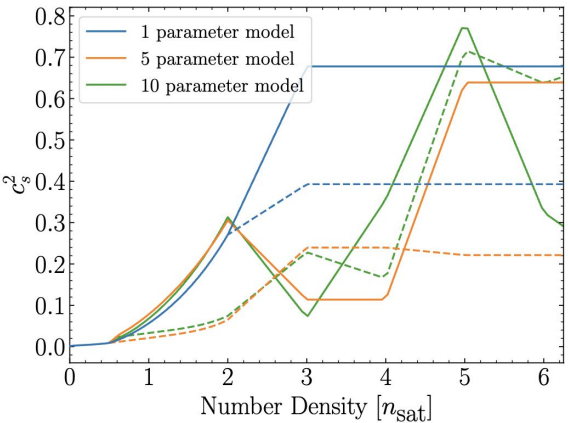
# Multilayer Perceptrons for the TOV equations

Multilayer Perceptrons (MLP) are the simplest, dense, feedforward neural networks. We use the method of deep ensembles where we use a set of 100 MLPs

Equation of State parameters

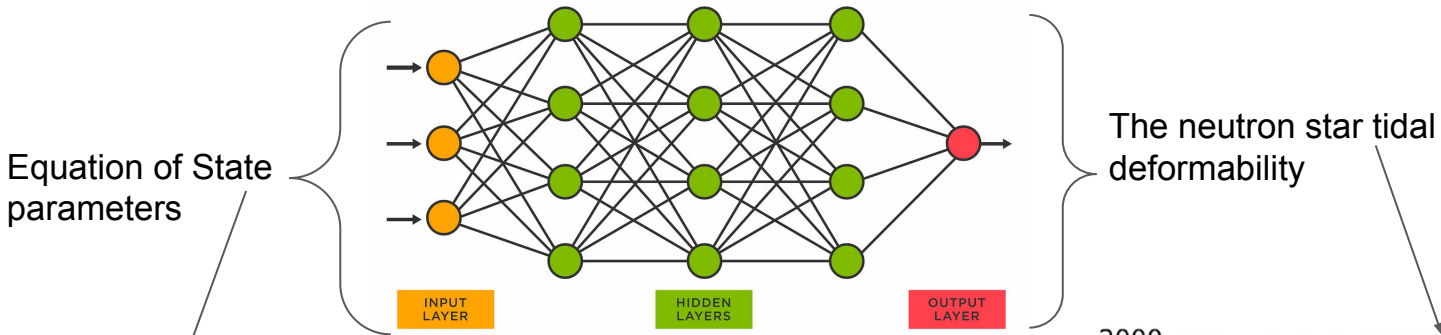


(x100)

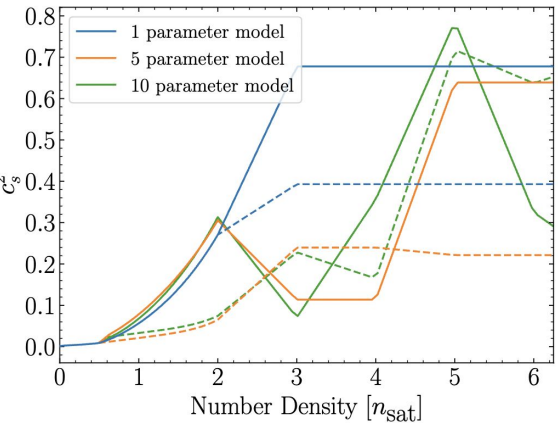


# Multilayer Perceptrons for the TOV equations

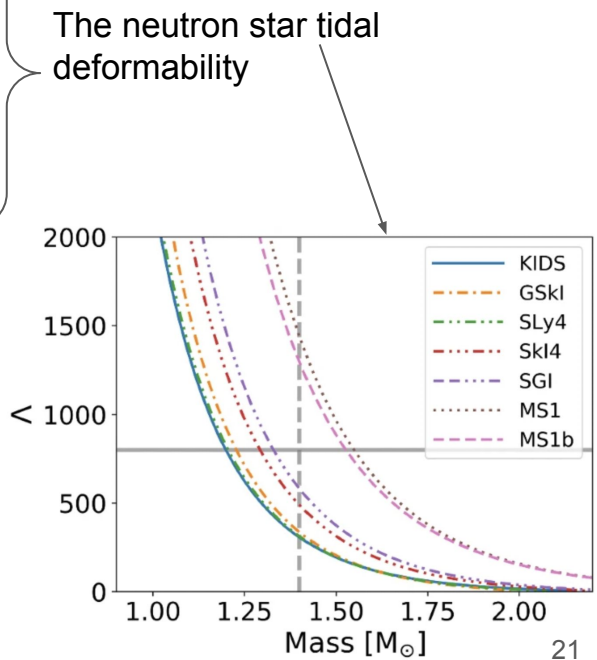
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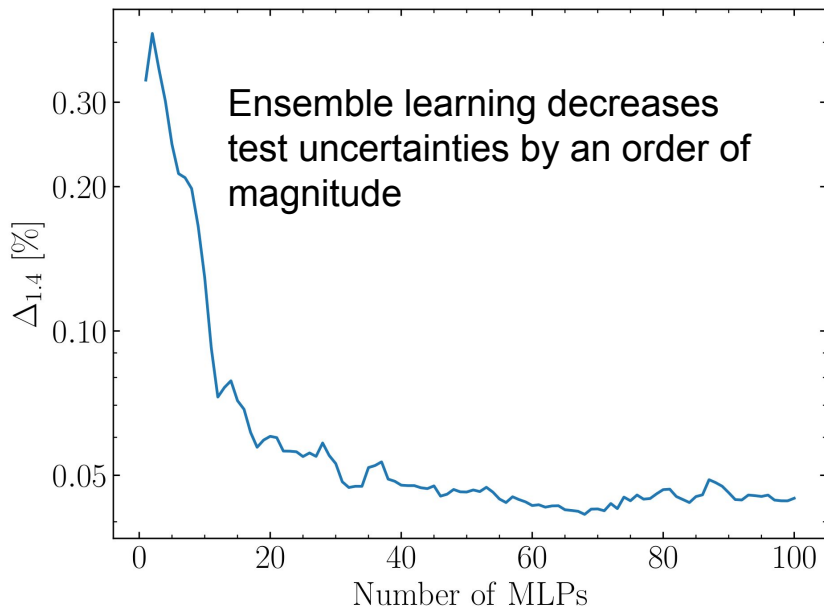
(x100)



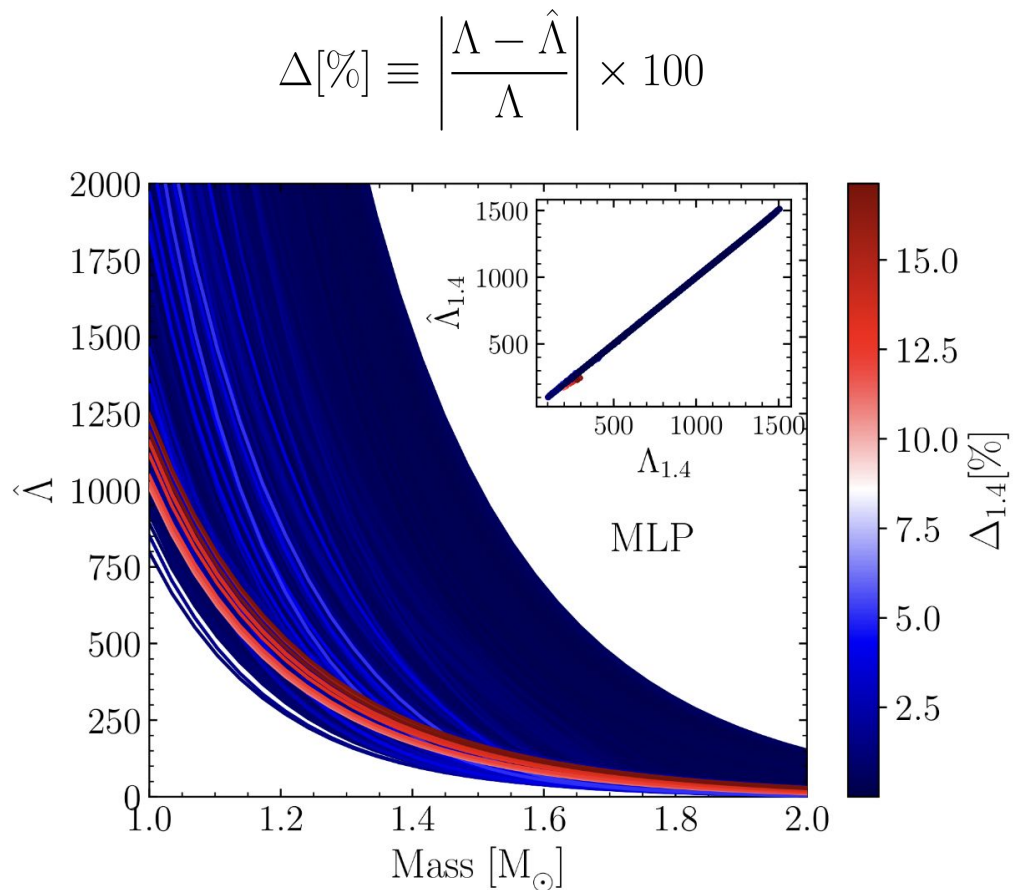
We generated 200K samples, randomly split into training and test data



# Multilayer Perceptrons for the TOV equations



- Average uncertainty on test samples is 0.04 %
- 0.01% of test samples are outliers (uncertainty > 10%)





# Summary

Speed-up factor  $\sim 10^7$   
 Uncertainty  $\sim 0.1\%$

Speed-up factor  $\sim 10^3$   
 Uncertainty  $\sim 0.04\%$



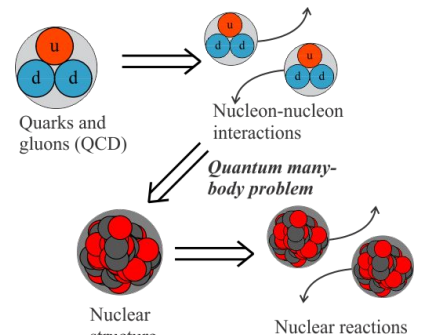
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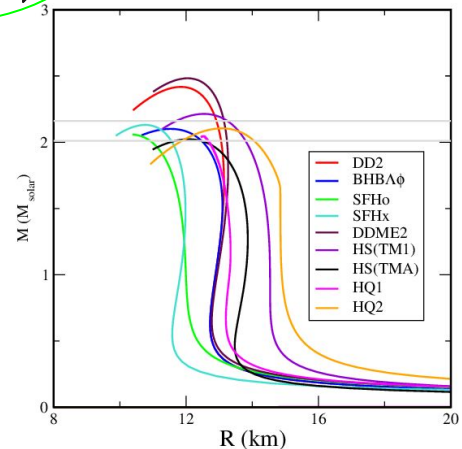
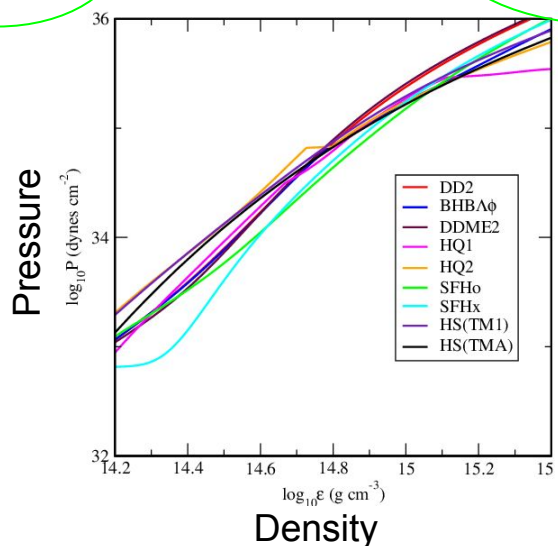
The Equation of State

$$G^{\mu\nu} = \kappa T^{\mu\nu}$$

Neutron star observables



$$V_{NN} = c_1(\Lambda) + \dots + c_j(\Lambda) + \dots + c_k(\Lambda) + \dots$$



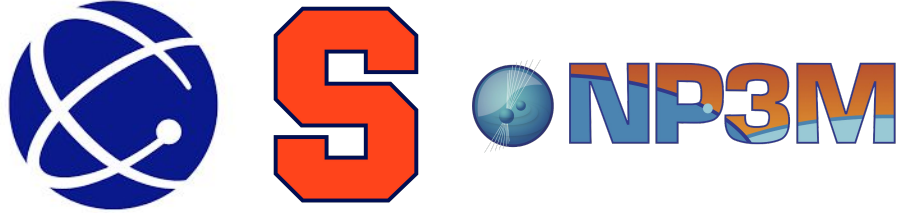
(In Progress)

Bayesian Inference methods

# Acknowledgements

Collaborators:

Ingo Tews, Duncan Brown, Achim Schwenk, Stefano Gandolfi, Collin Capano, Soumi De, Cassandra Armstrong, Brendan Reed, Pablo Giuliani, Kyle Godbey, Andrew Deneris

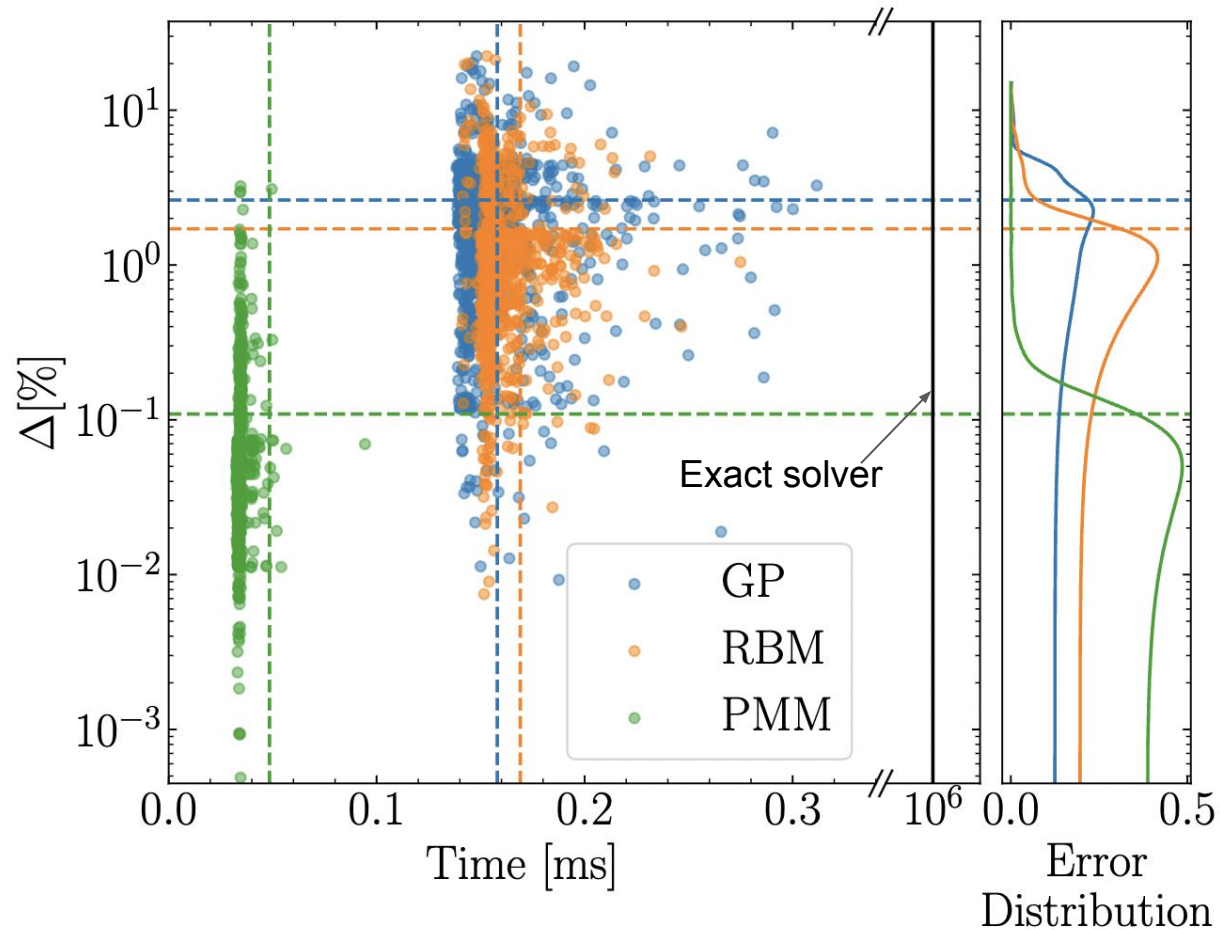


Thank You!



# Backup slides

# The trade-off between speed and computational accuracy



# The tradeoff between speed and accuracy

